

Back and forth from cool core to non-cool core: clues from radio-halos

M. Rossetti^{1,2}, D. Eckert¹, B. M. Cavalleri^{2,1}, S. Molendi¹, F. Gastaldello^{1,3}, and S. Ghizzardi¹

¹ Istituto di Astrofisica Spaziale e Fisica Cosmica, INAF, via Bassini 15, I-20133 Milano, Italy

² Dipartimento di Fisica – Università degli Studi di Milano via Celoria 16, Milano, Italy

³ University of California at Irvine, 4129, Frederick Reines Hall, Irvine, CA 92697-4575, USA

Preprint online version: June 24, 2011

Abstract

X-ray astronomers often divide galaxy clusters into two classes: “cool core” (CC) and “non-cool core” (NCC) objects. The origin of this dichotomy has been the subject of debate in recent years, between “evolutionary” models (where clusters can evolve from CC to NCC, mainly through mergers) and “primordial” models (where the state of the cluster is fixed “ab initio” by early mergers or pre-heating). We found that in a well-defined sample (clusters in the GMRT Radio halo survey with available *Chandra* or *XMM-Newton* data), none of the objects hosting a giant radio halo can be classified as a cool core. This result suggests that the main mechanisms which can start a large scale synchrotron emission (most likely mergers) are the same that can destroy CC and therefore strongly supports “evolutionary” models of the CC-NCC dichotomy. Moreover combining the number of objects in the CC and NCC state with the number of objects with and without a radio-halo, we estimated that the time scale over which a NCC cluster relaxes to the CC state, should be larger than the typical life-time of radio-halos and likely shorter than ≈ 3 Gyr. This suggests that NCC transform into CC more rapidly than predicted from the cooling time, which is about 10 Gyr in NCC systems, allowing the possibility of a cyclical evolution between the CC and NCC states.

Key words. Galaxies:clusters:general–X-rays:galaxies:clusters

1. Introduction

Galaxy clusters are often divided by X-ray astronomers into two classes: “cool core”(CC) and “non-cool core” (NCC) objects on the basis of the observational properties of their central regions. One of the open questions in the study of galaxy clusters concerns the origin of this distribution. The original model which prevailed for a long time assumed that the CC state was a sort of “natural state” for the clusters, and the observational features were explained within the old “cooling flow” model: radiation losses cause the gas in the centers of these clusters to cool and to flow inward. Clusters were supposed to live in this state until disturbed by a “merger”. Indeed, mergers are very energetic events that can shock-heat (Burns et al. 1997) and mix the ICM (Gómez et al. 2002); through these processes they were supposed to efficiently destroy cooling flows (e. g. McGlynn & Fabian 1984). After the mergers, clusters were supposed to relax and go back to the cooling flow state in a sort of cyclical evolution (e. g. Buote 2002). With the fall of the “cooling flow” model brought about by the *XMM-Newton* and *Chandra* observations (e. g. Peterson et al. 2001), doubts were cast also on the interpretation of mergers as the dominant mechanism which could transform CC clusters into NCC. These doubts were also motivated by the difficulties of numerical simulations in destroying simulated cool cores with mergers (e. g. Burns et al. 2008 and references therein). More generally speaking, the question arose whether the observed distribution of clusters was due to a primordial division into the two classes or rather to evolutionary differences during the history of the clusters.

For instance McCarthy et al. (2004, 2008) suggested that early episodes of non-gravitational pre-heating in the redshift range $1 < z < 2$ may have increased the entropy of the ICM of some

proto-clusters which did not have time to develop a full cool core. Burns et al. (2008) suggested that while mergers cannot destroy simulated cool cores in the local Universe, early major mergers could have destroyed nascent cool cores in an earlier phase of their formation ($z > 0.5$).

However, the “evolutionary” scenario, where recent and on-going mergers are responsible for the CC-NCC dichotomy, has been continuously supported by observations. Indeed, correlations have been found between the lack of a cool core and several multi-wavelength indicators of on-going dynamical activity (e. g. Sanderson et al. 2006, 2009, Leccardi et al. 2010).

Giant radio halos (RH) are the most spectacular evidence of non thermal emission in galaxy clusters (Ferrari et al. 2008 for a recent review). Over the last years, there has been increasing evidence in the literature that they are found in clusters with a strong on-going dynamical activity (e. g. Buote 2001; Schuecker et al. 2001, Hudson et al. 2010) suggesting that mergers could provide the energy necessary to accelerate (or re-accelerate) electrons to radio-emitting energies (Sarazin 2002; Brunetti et al. 2009). Recently, the connection between radio halos and mergers has been quantitatively confirmed on a well-defined statistical sample by Cassano et al. (2010).

In the framework of “evolutionary” scenarios, mergers are also responsible for the CC-NCC dichotomy. Therefore, we expect mergers to cause a relation between the absence of a cool core and the presence of a giant radio halo. The aim of this work is to assess statistically the presence of this relation on a well-defined sample and to test our “evolutionary” interpretation of the origin of the CC-NCC distribution.

Throughout this paper we assume a Λ CDM cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$ and $\Omega_\Lambda = 0.7$.

2. The sample

The choice of the sample is an important part of this project. We do not want to select an “archive radio sample” of clusters, since the search for radio halos often concentrated in clusters which showed some indications of a disturbed dynamical state from other wavelengths, such as the absence of a cool core in X-rays. Unlike archival samples, the “GMRT radio halo survey” (Venturi et al. 2007, 2008) is an excellent starting point for our aims: it consists of a deep pointed radio survey of a representative sample of 50 clusters selected in X-rays from flux-limited *ROSAT* surveys (REFLEX and eBCS), with $z = 0.2-0.4$, $L_X > 5 \times 10^{44} \text{ ergs s}^{-1}$ and $-30^\circ < \delta < 60^\circ$. For 35 clusters of this sample, uniformly observed with GMRT, Venturi et al. (2008) could either detect extended radio emission or put strong upper limits on it. The GMRT RH sample is designed to cover a well defined range in X-ray luminosity, such that according to the $P_{1.4\text{GHz}} - L_X$ relation (e. g. Brunetti et al. 2009), any radio-halo should have been detected in the survey. Therefore we can consider “radio-quiet” the clusters for which Venturi et al. (2008) did not find extended central radio emission.

We looked into the *Chandra* and *XMM-Newton* archives for observations of the clusters in the GMRT RH sample. We preferentially used *Chandra* observations to exploit the better angular resolution but we discarded observations with less than 1500 counts in each of the regions from which we extract spectra (see Sec. 3.3), moving to *XMM-Newton* when available.

To avoid confusion, we do not consider here the three objects in the GMRT RH survey with mini radio halos (A 2390, RXCJ 1504.1 – 0248 and Z 7160). Their central radio emission has been classified as a mini-halo because it extends for less than 500 kpc and it is associated to an active central radio galaxy (Bacchi et al. 2003; Mazzotta & Giacintucci 2008; Giacintucci et al. 2011). Mini-halos should not be confused with giant RHs since their origin is likely due to electrons injected by the central AGN re-accelerated by sloshing (Mazzotta & Giacintucci 2008; ZuHone et al. 2011) or AGN-driven turbulence (Cassano et al. 2008).

Our final sample consists of 22 clusters with available *Chandra* or *XMM-Newton* observations (Table 1). We verified this subsample to be representative of the starting sample (50 objects) in terms of luminosity with a Kolmogorov–Smirnov test. 10 clusters in our subsample are “radio-loud” (hosting a giant radio halo obeying the well known relation between the radio power at 1.4 GHz, $P_{1.4}$, and the X-ray luminosity L_X) and the remaining 12 are “radio-quiet”, well separated in the $P_{1.4} - L_X$ plane (see Brunetti et al. 2009 for a detailed discussion of this distribution). We note here that our sample of “radio-loud” clusters is composed of all the clusters with a confirmed giant radio halo in Venturi et al. (2008), with the addition of A697 and A1758 which were classified as “candidate halos” and were confirmed later (Macario et al. 2010; Giovannini et al. 2009).

We verified that in our sample clusters with radio-halos are neither more luminous in X-rays nor more distant than radio-quiet clusters.

3. Data analysis

3.1. Chandra data reduction

We analyzed *Chandra* data with the software CIAO v.4.1 and the calibration database (CALDB) 4.1.1. All data were reprocessed from the level 1 event files, following the standard *Chandra* data-

Table 1: Cool core indicators for the clusters of the sample.

Number	Cluster	$K_0(\text{keV cm}^2)$	σ
1	A 2163	437.98 ± 82.56	0.63 ± 0.03
2	A 521*	259.87 ± 36.25	0.57 ± 0.05
3	A 2219	411.57 ± 43.16	0.77 ± 0.06
4	A 2744	438.44 ± 58.71	0.90 ± 0.14
5	A 1758	230.84 ± 37.22	0.46 ± 0.04
6	RXCJ 2003.5 – 2323 [†]	340.61 ± 28.62	0.91 ± 0.10
7	A 1300*	97.26 ± 22.98	0.54 ± 0.06
8	A 773	244.32 ± 31.73	0.71 ± 0.08
9	A 697 [†]	178.09 ± 28.62	0.71 ± 0.10
10	A 209*	105.50 ± 26.94	0.58 ± 0.04
11	A 1423	68.32 ± 12.85	0.41 ± 0.05
12	A 2537	110.41 ± 19.37	0.36 ± 0.04
13	A 2631*	308.81 ± 37.38	0.8 ± 0.1
14	A 2667	19.32 ± 3.39	0.32 ± 0.02
15	A 3088	82.78 ± 8.42	0.30 ± 0.03
16	A 611	124.93 ± 18.61	0.47 ± 0.03
17	RXCJ 1115.8 + 0129	14.76 ± 3.09	0.26 ± 0.02
18	RXJ 1532.9 + 3021	16.93 ± 1.81	0.34 ± 0.03
19	RXJ 2228.6 + 2037*	118.79 ± 39.19	0.61 ± 0.06
20	Z 2701	39.66 ± 3.92	0.36 ± 0.01
21	Z 2089 [†]	14.59 ± 1.14	0.36 ± 0.02
22	AS 0780 [†]	22.24 ± 1.36	0.26 ± 0.01

Notes. The first column refers to the numbering of the objects in Fig. 1. Clusters 1-10 are “radio-halos” while objects 11-22 are “radio-quiet”.

(*) Pseudo-entropy ratio measured with *XMM-Newton* data. (†) *Chandra* observation not in the ACCEPT archive.

reduction threads¹. We applied the standard corrections to account for a time-dependent drift in the detector gain and charge transfer inefficiency, as implemented in the CIAO tools. From low surface brightness regions of the active chips we extracted a light-curve (5 – 10 keV) to identify and excise periods of enhanced background. We removed point sources detected with the CIAO tool *wavdetect*. Background analysis was performed using the blank-sky datasets provided in the CALDB. Background files were reprocessed and reprojected to match each observation. We extracted spectra from the background and source files from an external region not contaminated by cluster emission and we quantified the ratio between the count rate of the observation and of the background in a hard energy band (9.5 – 12 keV). By rescaling the background files for this number, we took into account possible temporal variations of the instrumental background dominating at high energies. This procedure does not introduce substantial distortions in the soft energy band, where the cosmic background components are more important, since we limit our analysis to regions where the source outshines the background in the soft band by more than one order of magnitude.

3.2. XMM-Newton data reduction

We have analyzed *XMM-Newton* observations for the 6 clusters (see Table 1) of our sample where we found a total number of *Chandra* counts < 1500 in the *IN* region (Sec. 3.3.2) used to measure the pseudo-entropy ratios (Leccardi et al. 2010). We generated calibrated event files using the SAS software v. 9.0 and then we removed soft proton flares using a double filtering process, first in a hard (10 – 12 keV) and then in a soft (2 – 5 keV) energy range. The event files were filtered according to PATTERN and FLAG criteria. Bright point-like sources were detected us-

¹ <http://cxc.harvard.edu/ciao/threads/index.html>

ing a procedure based on the SAS task `edetect_chain` and removed from the event files. As background files, we merged nine blank field observations, we reprojected them to match each observation and renormalized by the ratio of the count-rates in an external region as for the *Chandra* case (more details in Leccardi et al. 2010). We extracted spectra for the three EPIC detectors and fitted them separately. After verifying that best fit parameters from each instrument are consistent at less than 2σ , we combined them with a weighted mean.

3.3. Cool core indicators

For each of the clusters in our sample we have calculated two estimators of the core state: the central entropy K_0 (Cavagnolo et al. 2009, Sec. 3.3.1) and the pseudo-entropy ratio σ (Leccardi et al. 2010, Sec. 3.3.2). To measure them, both with *Chandra* and *XMM-Newton*, we extracted spectra and generated an effective area (ARF) for each region, which we associated to the appropriate response file (RMF). We fitted spectra within XSPEC v.12 with an absorbed *mekal* model, where we fixed redshift as given by NED² and N_H to the Dickey & Lockman (1990) values (for consistency with Cavagnolo et al. 2009). We verified that the N_H values agree with those derived from the LAB map (Kalberla et al. 2005) within 20% for all but one cluster of the sample. For this cluster, temperatures and normalizations obtained with the two different values of N_H are consistent at less than 1σ .

3.3.1. Central Entropy K_0

K_0 is derived from the fit of the entropy profile with the model $K_0 + K_{100}(r/100 \text{ kpc})^\alpha$. When available, we have used the values reported in the ACCEPT catalogue³ (Cavagnolo et al. 2009). For the 4 objects (see Table 1) whose *Chandra* observations were not public at the time of the compilation of ACCEPT, we have extracted the entropy profile following the same procedure as Cavagnolo et al. (2009). More specifically, we combined the temperature profile, measured by fitting a thermal model to spectra extracted in concentric annuli with at least 2500 counts, with the gas density profile. The latter was derived using the deprojection technique of Kriss et al. (1983) by combining the surface brightness profile with the spectroscopic count rate and normalization in each region of the spectral analysis. Errors were estimated with a Monte Carlo simulation.

3.3.2. Pseudo-entropy ratios σ

The pseudo-entropy ratio is defined as $\sigma = (T_{IN}/T_{OUT}) * (EM_{IN}/EM_{OUT})^{-1/3}$, where T is the temperature, EM is the emission measure (XSPEC normalization of the *mekal* model divided by the area of the region). The *IN* and *OUT* regions are defined with fixed fractions of R_{180} ($R < 0.05R_{180}$ for the *IN* region and $0.05R_{180} < R < 0.2R_{180}$ for the *OUT* region). We calculated R_{180} from T_{OUT} using the expression in Leccardi & Molendi (2008), iterating the process until it converged to stable values of R_{180} (Rossetti & Molendi 2010). The center from which we defined our region is the same used in ACCEPT (i. e. the X-ray peak or the centroid of the X-ray emission if these two points differ for more than 70 kpc).

The limited spatial resolution of EPIC may be an issue for measuring σ in clusters at $z > 0.25$, since it may cause the spreading

of photons coming from the *IN* circle into the *OUT* region (and vice-versa, although this contamination is likely to be less important especially in CC). Therefore, we adopted the cross-talk modification of the ARF generation software and fitted simultaneously the spectra of the two regions (see Snowden et al. 2008 and Ettori et al. 2010 for details).

4. Results

In Table 1, we report K_0 and σ for all the clusters of our sample. We classified clusters into core classes according to these indicators. For the sake of simplicity, we decided to classify them into two classes (CC and NCC) well aware of the existence of objects with intermediate properties (Leccardi et al. 2010).

Basing on K_0 , we divided the clusters population into CC ($K_0 < 50 \text{ keV cm}^2$) and NCC ($K_0 > 50 \text{ keV cm}^2$) as in Cavagnolo et al. (2009). Using this classification, we found that all “radio-loud” clusters are classified as NCC while “radio quiet” objects belong to both classes⁴ (Fig. 1 left panel). Because of the relatively low number of objects in our sample, we have to verify our result with Monte Carlo simulations to exclude that it comes out just from statistical fluctuations. To this aim, we have calculated the mean K_0 of our sample of radio loud clusters ($K_0 = 274 \pm 14 \text{ keV cm}^2$) and compared it with the distribution of the mean K_0 of 10 clusters randomly selected in the ACCEPT archive (Fig. 2 left panel). We found that the probability of finding by chance a mean K_0 larger than the value of the radio-loud sample is only 0.003% (0.002% – 0.007% considering the errors on the mean observed value). We have performed the same simulation randomly selecting clusters from the representative HIFUGS subsample (instead of the whole ACCEPT archive) finding even lower probabilities ($P = 10^{-4}\%$), as well as with the subsample of clusters in ACCEPT with redshift in the range $0.2 - 0.4$ ($P = 2 \cdot 10^{-4}\%$).

We have performed the same analysis using the pseudo-entropy ratios σ , using the threshold in Rossetti & Molendi (2010) to divide objects into classes (CC if $\sigma < 0.45$, NCC if $\sigma > 0.45$). Again, we found that none of the radio-loud clusters is classified as a CC while radio-quiet objects belong to both classes (Fig. 1 right panel). As for K_0 , we have performed a Monte Carlo simulation (Fig. 2), calculating the mean of our “radio-loud” sample ($\sigma = 0.69 \pm 0.02$) and comparing it with the distribution of the mean of 10 randomly selected values in the sample of Leccardi et al. (2010). We found a chance probability of finding a mean value larger than the observed value of 0.26% (0.02% – 1.96% if we consider the errors on the mean observed σ).

Finally, in Fig. 3 we plot the objects of our sample in the $\sigma - K_0$ plane. No radio-loud cluster is found in the lower left quadrant of the plot, that we define by subtracting the corresponding error from the minimum K_0 (σ) value of the clusters in the radio-loud sample (dashed lines in Fig. 3). We run 10^5 Monte-Carlo simulations in the $\sigma - K_0$ plane, allowing the points to vary within the error bars and randomly selecting 10 objects between our simulated points. In only 2 out of 10^5 simulations, no cluster is found in the lower left quadrant of the plot, proving that our result is statistically significant at 4.3σ .

5. Discussion

As discussed in Sec. 4, we found a robust correlation between the presence of a giant radio halo and the absence of a cool core, as

² <http://nedwww.ipac.caltech.edu/>

³ <http://www.pa.msu.edu/astro/MC2/accept/>

⁴ A qualitatively similar result was also reported by Enßlin et al. (2011).

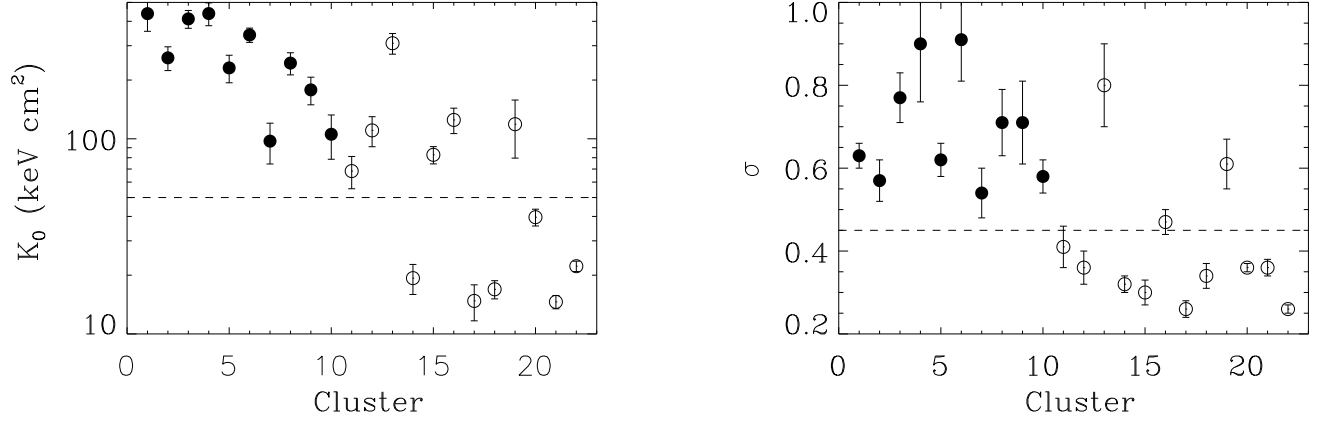


Figure 1: Cool core indicators (K_0 and σ) for all the clusters in the sample. Filled symbols are “radio-halo” clusters while open symbols are “radio-quiet”. Dashed lines mark the threshold between CC (below) and NCC (above). The numbering on the “x” axis refers to the position of the objects in Table 1. Error bars are 1σ .

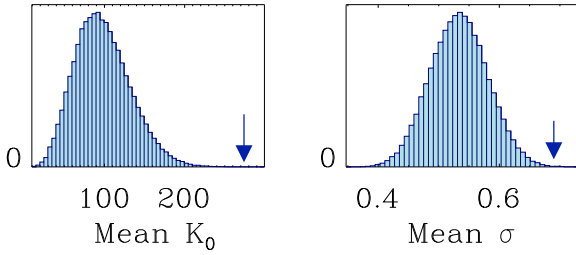


Figure 2: Histogram density for the mean K_0 (left) and σ (right) of 10 randomly selected clusters in 10^5 Monte Carlo simulations. The arrows show the observed mean values in our radio-halos sample.

Table 2: Statistical results of Monte Carlo simulation.

Test	Value	Null hypothesis probability
Mean K_0	274 keV cm ²	0.003% (4.17 σ)
	260 – 288 keV cm ²	0.007 – 0.002% (3.98 – 4.26 σ)
Mean σ	0.69	0.26%, (3.01 σ) [†]
	0.67 – 0.71	1.96 – 0.02% (2.33 – 3.71 σ)
σ vs K_0		0.002% (4.26 σ)

Notes. ^(†) This is the value we will refer to in the text.

indicated by both K_0 and σ . Despite the relatively low number of objects considered this result is statistically significant as shown by the outputs of our Monte Carlo simulations summarized in Table 2. Indeed, we found that the probability of a chance result is almost always negligible and lower than 2% even in the worst case. Moreover, we have obtained these results on a well defined sample that unlike archival samples is not biased towards clusters with RHs and without CC (Sec. 2).

One may argue that our choice of excluding the mini-halos clusters from this sample may hamper our results, since this could

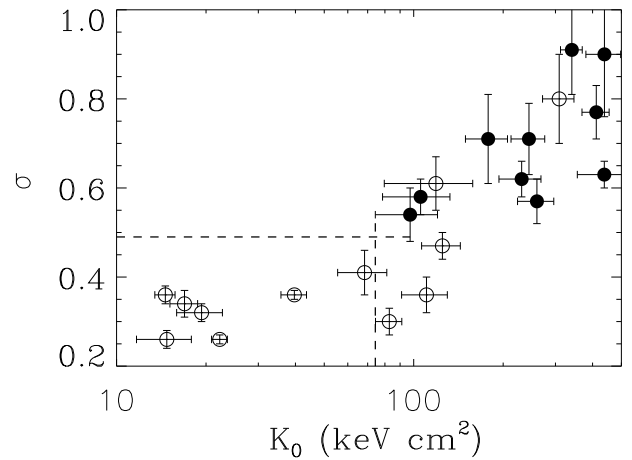


Figure 3: Central entropy and pseudo-entropy ratio for radio-loud (filled circles) and radio-quiet (open circles) clusters. The dashed lines indicate the quadrant of the plot where no-radio halo is found.

be exactly the clusters with both a RH and a CC⁵. We emphasize here that the classification of these objects as mini-halos was based only on the properties of their radio-emission (e.g. their sizes), regardless of the X-ray properties of their host clusters. Mini-radio halos are still poorly understood sources (e.g. Murgia et al. 2009), thus we cannot exclude that they could be a phase during the evolution of giant RHs. However, since they are usually considered as a different class of objects from giant RHs with a different origin (Mazzotta & Giacintucci 2008; Cassano et al. 2008; ZuHone et al. 2011), in the following we work under this hypothesis.

The results presented in this paper have important implications

⁵ We quote here the value of K_0 in ACCEPT for the three clusters with mini radio-halos. A2390: 14.73 ± 6.99 , RXCJ1504.1–0248: 13.08 ± 0.94 and Z7160: 16.88 ± 1.52 (Cavagnolo et al. 2009)

on the origin of the CC-NCC dichotomy (Sec. 5.1), since they suggest that the processes leading to the formation of RHs are likely the same that destroy CCs. They also give us the opportunity to estimate the time-scale over which a NCC cluster can relax to the CC state (Sec. 5.2).

5.1. Origin of the CC-NCC dichotomy.

The result presented in this paper are naturally addressed in “evolutionary” scenarios of the CC-NCC dichotomy where recent and on-going mergers are responsible for the disruption of the cool cores and for the formation of radio halos. Conversely, alternative “primordial” scenarios would have to explain why radio-halos are found only in NCC object. Radio-halos are transient phenomena with a typical life time of ~ 1 Gyr (Brunetti et al. 2009) associated therefore to recent mergers. In the “primordial” model of Burns et al. (2008) early major mergers (at $z > 0.5$) destroy nascent cool cores and are responsible for the NCC clusters we observe today. However these early mergers cannot explain the radio emission in most of the radio loud clusters of our sample which have $z \sim 0.2$, corresponding to more than 2.5 Gyr from $z = 0.5$. Even for the most distant cluster of our sample (RXCJ2003.5 – 2323 with $z = 0.317$) the life time of the radio halo should be at least ~ 1.5 Gyr to reconcile it with the model of Burns et al. (2008). Therefore it is hard to explain the results of the present paper within the “primordial” model of Burns et al. (2008). It is even harder and against Occam’s razor to explain it in the frame-work of “primordial” pre-heating models (e. g. McCarthy et al. 2004) which require additional physical processes, completely unrelated to those responsible for the radio emission, to account for NCC clusters.

5.2. Implications on the relaxation time-scale

One of the major open issues in the “evolutionary” scenario of the CC-NCC dichotomy is the estimate of the likelihood of NCC systems to be transformed into CC objects. The typical cooling times⁶ in the central regions of NCC are larger than 10 Gyr (e. g. Rossetti & Molendi 2010), seeming to imply that once a system has been heated to a NCC state it will not revert to a CC in less than a Hubble time. Indeed, the mean cooling time of the NCC objects ($K_0 > 50 \text{ keV cm}^2$) in our sample is ≈ 15 Gyr.

We can use the results presented in this paper to estimate the ratio of the two relevant time-scales: the life time of radio-halos (t_{RH}) and the time-scale over which NCC relax to the CC state ($t_{NCC \rightarrow CC}$). The absence of clusters with RH and CC implies that the two processes leading to the formation of these objects (i. e. mergers creating a radio halo but preserving the CC and a fast relaxation from NCC to CC) are extremely unlikely. Indeed the fact that we do not observe this class of clusters but we do observe many NCC objects without a radio halo, implies that RH clusters lose their radio emission more rapidly than developing a new CC, thus $t_{NCC \rightarrow CC} > t_{RH}$. It is possible to show that in stationary conditions this also implies that we cannot have mergers creating a radio halo but preserving the CC (see Appendix A).

In this framework, we can provide also an upper limit to the ratio of the two time-scales, which depends on the ratio of the num-

ber of radio-quiet clusters with a CC ($N_{RQ,CC}$) to those which are NCC ($N_{RQ,NCC}$) through the expression:

$$1 < \frac{t_{NCC \rightarrow CC}}{t_{RH}} \leq \frac{N_{RQ,NCC}}{N_{RQ,CC}} \frac{N_{RQ}}{N_{RH}} \quad (1)$$

where N_{RQ} and N_{RH} are the total number of radio-quiet clusters and of radio-halos respectively (see Appendix A for the derivation of this expression). The second fraction in the right term of expression 1 is one of the main results of the GMRT RH survey: $N_{RQ}/N_{RH} = 25/10$, where we have considered as radio-quiet also the clusters with mini-halos and/or relics, as discussed in Sec. 2. Since 10 of the radio-quiet objects in the GMRT RH survey have not been observed by either *XMM-Newton* or *Chandra*, we cannot assess their core state and therefore we cannot really measure the value of the first fraction in the right term of expression 1. If we use K_0 as an indicator of the core state the value we measure in our subsample is $N_{RQ,NCC}/N_{RQ,CC} = 6/6$ and $6/9$ if we consider also the three mini-halos clusters which are all classified as CC. Considering the 10 clusters without X-ray observations this ratio could range from $6/19$ to $16/9$. However it is unlikely that the 10 unobserved clusters are either all CC or all NCC. Estimates of the fraction of CC objects in representative samples depend strongly on the indicator used to classify them (Chen et al. 2007; Hudson et al. 2010). If we consider the HIFLUGS subsample in the ACCEPT catalogue, the fraction of clusters with $K_0 < 50 \text{ keV cm}^2$ is 0.44. Given that, we expect 6 out of 10 missing clusters to be CC, corresponding to $N_{RQ,NCC}/N_{RQ,CC} = 10/15$. Allowing the CC fraction to be in the range 0.35–0.5 we expect from 3 to 8 CC in the unobserved clusters ($N_{RQ,NCC}/N_{RQ,CC} = 8/17$ – $13/12$). Assuming $t_{RH} = 1$ Gyr (Brunetti et al. 2009) we find an upper limit $t_{NCC \rightarrow CC} = 1.7$ Gyr with the observed ratio $6/9$ and with the expected ratio $10/15$, while in the range 1.2–2.7 Gyr if we allow the ratio to vary between $8/17$ and $13/12$.

While we cannot provide more precise constraints because not all clusters in the GMRT RH survey have been observed by *XMM-Newton* or *Chandra*, our results show that even in the unlikely case where all the unobserved clusters are NCC, the time scale over which a NCC cluster relaxes to the CC state is less than 4.5 Gyr. This time scale is significantly shorter than the typical cooling time of NCC clusters ≈ 10 Gyr (e. g. Rossetti & Molendi 2010).

We can predict that the ratio $N_{RQ,NCC}/N_{RQ,CC}$ should be 4 in the case $t_{NCC \rightarrow CC} = t_{cool} \approx 10$ Gyr and estimate with a Monte-Carlo simulation that this is inconsistent at more than 3.4σ with the permitted values in the GMRT sample assuming a CC fraction in the range 0.35–0.5.

As discussed in Rossetti & Molendi (2010), the cooling time derived from thermodynamical quantities should be considered as an upper limit for the relaxation time-scale. If, during the merger which destroyed the cool core, the mixing of the gas has not been completely effective and the ICM retains a certain degree of multiphaseness, then the cooler and denser phases will rapidly sink toward the center and re-establish a CC on time-scales shorter than the cooling time calculated under the assumption of uniform temperature and density. Current CCD instruments do not allow us to distinguish a multi-temperature structure with $kT > 2 \text{ keV}$ (Mazzotta et al. 2004) and this measurement will become possible only with the high spectral resolution imaging calorimeters on board ASTRO-H. In the meantime, the method we described above provides the opportunity to measure the dynamical time scale $t_{NCC \rightarrow CC}$ and thus to test the possibility of a “return journey” to the CC state.

⁶ We refer to the isobaric cooling time, see discussion in Peterson & Fabian (2006).

6. Conclusions

In this paper, we have analyzed the X-ray observations of the clusters in the GMRT RH survey (Venturi et al. 2007, 2008), finding that all clusters which host a radio-halo are also classified as NCC. Although obtained with a relatively low number of objects, this result is statistically significant at more than 3σ (Table 2).

This result implies that the mechanisms which generate radio-halos (most likely mergers) are the same that can destroy cool cores, supporting the “evolutionary” origin of the CC-NCC dichotomy. Moreover, we have shown that combining the number of radio-quiet and radio-halos objects with the number of CC and NCC, it is possible to provide upper and lower limits to the ratio of the two relevant time scales: the life time of the radio halo and the relaxation time from NCC to CC. Assuming $t_{RH} = 1$ Gyr (Brunetti et al. 2009), we constrained $t_{NCC \rightarrow CC}$ in the interval $1 - 2.7$ Gyr. These values are significantly shorter than the typical cooling time of NCC objects ($t_{cool} \approx 10$ Gyr, Rossetti & Molendi 2010), which predicted that most NCC would not develop a new cool-core in less than a Hubble time. On the contrary, the dynamical time-scale we have estimated in this paper allows a “return journey” to the CC state and suggests that the gas in the central regions of NCC clusters should be multi-phase. Only with the imaging calorimeters on board ASTRO-H will it be possible to test this important prediction, which may have strong implications on the physics of the ICM.

Acknowledgements. We would like to thank the referee T. Reiprich for his useful and valuable suggestions. We wish to thank P. Humphrey for the use of his *Chandra* code and K. Cavagnolo for the ACCEPT catalogue. MR acknowledges stimulating discussions with R. Cassano, G. Brunetti and T. Venturi. We are supported by ASI-INAF grant I/009/10/0, DE is supported also by the Occhialini fellowship at IASF-Milano.

References

- Bacchi, M., Feretti, L., Giovannini, G., & Govoni, F. 2003, *A&A*, 400, 465
 Brunetti, G., Cassano, R., Dolag, K., & Setti, G. 2009, *A&A*, 507, 661
 Buote, D. A. 2001, *ApJ*, 553, L15
 Buote, D. A. 2002, in *ASSL Vol. 272: Merging Processes in Galaxy Clusters*, 79–107
 Burns, J. O., Hallman, E. J., Gantner, B., Motl, P. M., & Norman, M. L. 2008, *ApJ*, 675, 1125
 Burns, J. O., Loken, C., Gomez, P., et al. 1997, in *Astronomical Society of the Pacific Conference Series*, Vol. 115, *Galactic Cluster Cooling Flows*, ed. N. Soker, 21–+
 Cassano, R., Ettori, S., Giacintucci, S., et al. 2010, *ApJ*, 721, L82
 Cassano, R., Gitti, M., & Brunetti, G. 2008, *A&A*, 486, L31
 Cavagnolo, K. W., Donahue, M., Voit, G. M., & Sun, M. 2009, *ApJS*, 182, 12
 Chen, Y., Reiprich, T. H., Böhringer, H., Ikebe, Y., & Zhang, Y.-Y. 2007, *A&A*, 466, 805
 Cohn, J. D. & White, M. 2005, *Astroparticle Physics*, 24, 316
 Dickey, J. M. & Lockman, F. J. 1990, *ARA&A*, 28, 215
 Eckert, D., Molendi, S., & Paltani, S. 2011, *A&A*, 526, A79+
 Enßlin, T., Pfrommer, C., Miniati, F., & Subramanian, K. 2011, *A&A*, 527, A99+
 Ettori, S., Gastaldello, F., Leccardi, A., et al. 2010, *A&A*, 524, A68+
 Ferrari, C., Govoni, F., Schindler, S., Bykov, A. M., & Rephaeli, Y. 2008, *Space Science Reviews*, 134, 93
 Giacintucci, S., Markevitch, M., Brunetti, G., Cassano, R., & Venturi, T. 2011, *A&A*, 525, L10+
 Giovannini, G., Bonafede, A., Feretti, L., et al. 2009, *A&A*, 507, 1257
 Gómez, P. L., Loken, C., Roettiger, K., & Burns, J. O. 2002, *ApJ*, 569, 122
 Hudson, D. S., Mittal, R., Reiprich, T. H., et al. 2010, *A&A*, 513, A37+
 Kalberla, P. M. W., Burton, W. B., Hartmann, D., et al. 2005, *A&A*, 440, 775
 Kriss, G. A., Cioffi, D. F., & Canizares, C. R. 1983, *ApJ*, 272, 439
 Leccardi, A. & Molendi, S. 2008, *A&A*, 486, 359
 Leccardi, A., Rossetti, M., & Molendi, S. 2010, *A&A*, 510, A82+
 Macario, G., Venturi, T., Brunetti, G., et al. 2010, *A&A*, 517, A43+
 Mazzotta, P. & Giacintucci, S. 2008, *ApJ*, 675, L9
 Mazzotta, P., Rasia, E., Moscardini, L., & Tormen, G. 2004, *MNRAS*, 354, 10

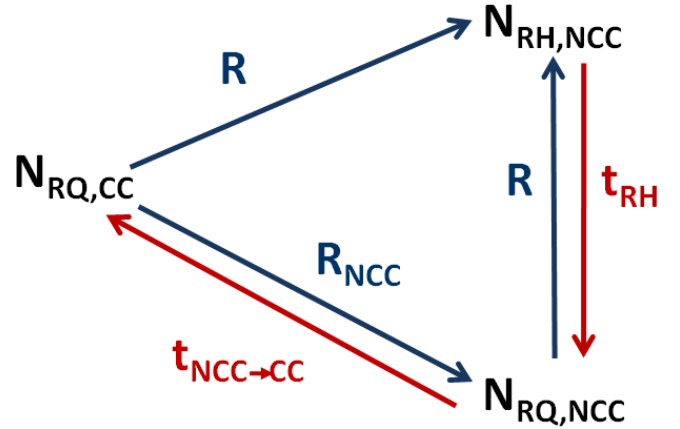


Figure A.1: A schematic view of the evolution of the thermodynamical and radio state in a cyclical evolutionary scenario. Mergers (blue arrows) are the physical processes that can create a radio-halo or destroy the cool core, while time (red arrows) is responsible for the fading of radio-halos (t_{RH}) and for the formation of CCs ($t_{NCC \rightarrow CC}$)

- McCarthy, I. G., Babul, A., Bower, R. G., & Balogh, M. L. 2008, *MNRAS*, 386, 1309
 McCarthy, I. G., Balogh, M. L., Babul, A., Poole, G. B., & Horner, D. J. 2004, *ApJ*, 613, 811
 McGlynn, T. A. & Fabian, A. C. 1984, *MNRAS*, 208, 709
 Murgia, M., Govoni, F., Markevitch, M., et al. 2009, *A&A*, 499, 679
 Peterson, J. R. & Fabian, A. C. 2006, *Phys. Rep.*, 427, 1
 Peterson, J. R., Paerels, F. B. S., Kaastra, J. S., et al. 2001, *A&A*, 365, L 104
 Rossetti, M. & Molendi, S. 2010, *A&A*, 510, A83+
 Sanderson, A. J. R., Edge, A. C., & Smith, G. P. 2009, *MNRAS*, 398, 1698
 Sanderson, A. J. R., Ponman, T. J., & O’Sullivan, E. 2006, *MNRAS*, 372, 1496
 Sarazin, C. L. 2002, in *ASSL Vol. 272: Merging Processes in Galaxy Clusters*, 1–38
 Schuecker, P., Böhringer, H., Reiprich, T. H., & Feretti, L. 2001, *A&A*, 378, 408
 Snowden, S. L., Mushotzky, R. F., Kuntz, K. D., & Davis, D. S. 2008, *A&A*, 478, 615
 Venturi, T., Giacintucci, S., Brunetti, G., et al. 2007, *A&A*, 463, 937
 Venturi, T., Giacintucci, S., Dallacasa, D., et al. 2008, *A&A*, 484, 327
 ZuHone, J., Markevitch, M., & Brunetti, G. 2011, in *Proceedings of the conference “Non-thermal phenomena in colliding galaxy clusters”* (Nice, 15-18 November 2010), astro-ph:1101.4627

Appendix A: Derivation of the time-scale $t_{NCC \rightarrow CC}$

The main result of our paper is that there are no RH clusters with a CC. Therefore we propose a simplified evolutionary scenario (shown in Fig. A.1) where clusters can be found in only three possible states: RH clusters which are also NCC ($N_{RH,NCC}$) and radio-quiet clusters which can be either CC ($N_{RQ,CC}$) or NCC ($N_{RQ,NCC}$). We assume that all the mergers that create a RH also destroy the cool core and they happen with a rate R . In addition there could be some mergers which destroy the cool core but do not create a RH, happening with a rate R_{NCC} . We denote the life-time of the radio-halo as t_{RH} and the time required for a NCC object to relax to the CC state as $t_{NCC \rightarrow CC}$. The underlying assumption (see Sec. 5.2) on the time-scales is that $t_{NCC \rightarrow CC} > t_{RH}$, thus clusters with radio halos first lose their radio emission and then develop a new CC.

In this scenario, we can write a system of continuity equations for the three states (see Eckert et al. 2011 for a similar argument):

$$\begin{cases} \frac{dN_{RQ,CC}}{dt} = \frac{N_{RQ,NCC}}{t_{NCC \rightarrow CC}} - (R + R_{NCC})N_{RQ,CC} \\ \frac{dN_{RH,NCC}}{dt} = R(N_{RQ,CC} + N_{RQ,NCC}) - \frac{N_{RH,NCC}}{t_{RH}} \\ \frac{dN_{RQ,NCC}}{dt} = \frac{N_{RH,NCC}}{t_{RH}} + R_{NCC}N_{RQ,CC} - \frac{N_{RQ,NCC}}{t_{NCC \rightarrow CC}} - RN_{RQ,NCC} \end{cases} \quad (A.1)$$

Any of these equations is not independent from the other two ($N_{RQ,CC} + N_{RQ,NCC} + N_{RH,NCC} = N_{clusters}$), therefore we decided to keep only the first two equations of the system A.1. We then make the further assumption of a stationary situation, which allows us to cancel the left terms in expression A.1, and we find:

$$\begin{cases} \frac{N_{RQ,NCC}}{t_{NCC \rightarrow CC}} = (R + R_{NCC})N_{RQ,CC} \\ \frac{N_{RH,NCC}}{t_{RH}} = RN_{RQ} \end{cases}, \quad (A.2)$$

where we used $N_{RQ,NCC} + N_{RQ,CC} = N_{RQ}$. Dividing the first equation by the second, we get

$$\frac{N_{RQ,NCC}}{N_{RH,NCC}} = \frac{t_{NCC \rightarrow CC}}{t_{RH}} \left(1 + \frac{R_{NCC}}{R}\right) \frac{N_{RQ,CC}}{N_{RQ}}. \quad (A.3)$$

Since $R_{NCC}/R \geq 0$,

$$\frac{N_{RQ,NCC}}{N_{RH,NCC}} \geq \frac{t_{NCC \rightarrow CC}}{t_{RH}} \frac{N_{RQ,CC}}{N_{RQ}}. \quad (A.4)$$

It is interesting to note that the inequality in the expression A.4 becomes an equation when $R_{NCC} = 0$, i. e. if merger events capable of destroying CC always generate radio halos.

In principle, we should have considered a four-state system allowing the possibility to have clusters with RH and a CC. We can write the continuity equation for this state, with the usual assumption of a stationary situation:

$$R_{RH,CC}N_{RQ,CC} = R_{NCC}N_{RH,CC} + \frac{N_{RH,CC}}{t_{RH}}, \quad (A.5)$$

where $R_{RH,CC}$ is the occurrence of mergers which can produce a RH without destroying the CC. However, the fact that $N_{RH,CC} = 0$, directly implies that $R_{RH,CC} = 0$, i. e. all the mergers capable of generating a RH also destroy the CC, thus justifying the construction of a three-state system such as in Fig. A.1

One may argue that the key assumption of a stationary situation to derive expression A.2, is not justified. Indeed the fact that the time scales involved are comparable to the time interval we are considering should prevent us from assuming equilibrium. However, we recall here that in this simplified system we are looking at a “snapshot” of an evolutionary process which happens on time scales of the order of some Gyrs. Unless the merger rate, which is the ultimate factor in determining the changes of state, were to vary abruptly between $z = 0.4$ and $z = 0.2$, our system should not be very far from equilibrium. Indeed numerical simulations (e. g. Cohn & White 2005) show that the merger rate should vary only smoothly with time. Moreover, it is possible to take into account the left terms in expression A.1 and show that if they are smaller (of the order of one tenth) with respect to the change in the number of massive clusters from their formation epoch, our results still hold. Furthermore, the scenario described here is very simplified also because it assumes that the relevant time scales are the same for all objects while it is more reasonable to expect a distribution of values. If the distributions of $t_{NCC \rightarrow CC}$ and t_{RH} overlap we could have some clusters for which $t_{NCC \rightarrow CC} < t_{RH}$ and therefore we could in principle

observe some clusters with both a RH and a CC. Addressing both these issues (the non-stationary situation and the distribution of the time-scales) is not possible with present data, since it requires to follow the evolution of a large sample of clusters over several Gyrs. However, the necessary observations will likely become possible in the next years, thanks to LOFAR and e-ROSITA.